

# Thresholds for the Dust Driven Mass Loss from C-rich AGB Stars

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**Abstract.** It is well established that mass loss from AGB stars due to dust driven winds cannot be arbitrarily low. We model the mass loss from carbon rich AGB stars using detailed frequency-dependent radiation hydrodynamics including dust formation. We present a study of the thresholds for the mass loss rate as a function of stellar parameters based on a subset of a larger grid of such models and compare these results to previous theoretical work. Furthermore, we demonstrate the impact of the pulsation mechanism and dust formation for the creation of a stellar wind and how it affects these thresholds and briefly discuss the consequences for stellar evolution.

**Keywords.** stars: AGB and post-AGB, atmospheres, mass loss

## 1. Why a Mass Loss Threshold?

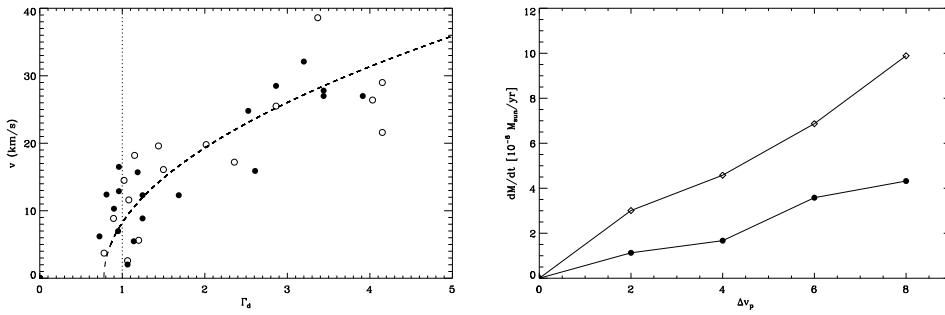
As shown by, e.g., Gail & Sedlmayr (1987), Dominik et al. (1990) and apparent in the detailed models by Höfner et al. (2003) as well, a dust-driven stellar wind cannot be maintained down to arbitrarily small ratio of radiative to gravitational acceleration  $\Gamma_d$ . For a "polytropic wind" one may derive an analytical expression for the terminal wind velocity,

$$v_\infty^2 \approx \frac{1}{2} \Delta v_p^2 + \left( \frac{2}{\gamma - 1} \right) \bar{c}_s^2(R_{in}) + \bar{v}_{esc}^2(R_c) \left[ \Gamma_d - \frac{R_c}{R_{in}} \right], \quad (1.1)$$

where  $\gamma$  is the polytropic index,  $R_c$  is the characteristic radius at which dust starts to condense,  $\bar{v}_{esc}^2(R_c)$  is the average escape velocity at  $R_c$ ,  $\bar{c}_s^2(R_{in})$  is the average sound speed at the inner boundary of the model (located at  $R_{in} \sim R_\star$ ) and  $\Delta v_p$  is the "piston amplitude", i.e. the strength of the pulsations (Mattsson 2006). The equation above captures the general trend of  $v_\infty$  with  $\Gamma_d$  and predicts a threshold at  $\Gamma_d \approx 0.8$  for reasonable values of the model parameters (see Fig. 1, left panel).

We have used our RHD code for dynamic stellar atmospheres of carbon-rich AGB stars (described in Höfner et al. 2003, Mattsson 2006), including frequency-dependent radiative transfer and dust formation, to explore the relations between basic stellar parameters and a dust-driven stellar wind. Here we present results from the computation of a grid of wind models at solar metallicity. An associated library of dynamic spectra is under development (see the poster by Wahlin et al.).

A mass loss threshold appears as one would expect from Eq. (1.1) and we find that below a critical C/O and/or above a critical  $T_{eff}$  no dust driven wind can be formed. All other stellar parameters were kept fixed in these models. We also see a rather strong dependence on C/O for both the wind velocity and the mass loss rate, which is quite interesting in comparison with previous studies of this kind. Arndt et al. (1997) as well as Wachter et al. (2002) find a weak dependence on C/O, which stands in sharp contrast to the results presented here. However, our findings here (as well as in Höfner et al. 2003) are, qualitatively speaking, hardly a new discovery. Höfner and Dorfi (1997) and Winters



**Figure 1.** Left: Wind velocity (in  $\text{km s}^{-1}$ ) as function of the acceleration parameter  $\Gamma_d$  for all models of the sub-grid calculated with  $M_\star = 1M_\odot$ ,  $Z = Z_\odot$  and  $\Delta v_p = 4.0 \text{ km s}^{-1}$ . Black dots represent models with  $\Delta v_p = 4.0 \text{ km s}^{-1}$  and circles represent models with  $\Delta v_p = 6.0 \text{ km s}^{-1}$ . The dashed curve shows an analytical model with  $\gamma = 7/6$ ,  $M_\star = 1M_\odot$ ,  $R_\star = 3.5 \cdot 10^{13} \text{ cm}$ ,  $R_{\text{in}}(0) = 0.9R_\star$ ,  $R_c = 2.5R_\star$  and  $c_s(R_{\text{in}}) = 7.0 \text{ km s}^{-1}$ . Right: The mass loss rate as a function of the piston amplitude. Black dots represents the case where  $L_\star = 7100L_\odot$  and diamonds represents  $L_\star = 10000L_\odot$ .

et al. (2000) have already pointed out the strong C/O-dependence, especially in the critical wind regime, although this has not been widely recognised. Furthermore, there is a linear dependence of the mass loss rate on the piston amplitude, i.e.  $\dot{M} \propto \Delta v_p$  (see Fig. 1, right panel). The trend is strong enough to make  $\Delta v_p$  significant in parametric prescriptions of the mass loss rate.

## 2. Conclusions

The results from our new detailed grid of wind models at solar metallicity suggests that C-stars with strong winds may actually be a rare species. How would this affect models of stellar evolution, nucleosynthesis and, consequently, models of chemical evolution of galaxies? We want to make the following points:

- The strength of the pulsations and the C/O-ratio are *not* redundant parameters in a mass loss prescription.
- It may be dangerous to use parametric mass loss formulae including too few stellar parameters and extrapolate beyond the range of stellar parameters used obtain the formula.
- The mass loss rate depends strongly on the efficiency of dust formation, which cannot simply be parameterised in terms of the basic stellar parameters: mass, luminosity and temperature, only.
- **There exists a threshold for dust-driven winds which has previously been neglected in mass loss prescriptions and thus not included in models of stellar evolution!**

## References

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